

LETTER TO THE EDITOR

Planck's confirmation of the M31 disk and halo rotation

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ABSTRACT

Planck's data acquired during the first 15.4 months of observations towards both the disk and halo of the M31 galaxy are analyzed. We confirm the existence of a temperature asymmetry, previously detected by using the 7-year WMAP data, along the direction of the M31 rotation, therefore indicative of a Doppler-induced effect. The asymmetry extends up to about 10^0 ($\simeq 130$ kpc) from the M31 center. We also investigate the recent issue raised in Rubin and Loeb (2014) about the kinetic Sunyaev-Zeldovich effect from the diffuse hot gas in the Local Group, predicted to generate a hot spot of a few degrees size in the CMB maps in the direction of M31, where the free electron optical depth gets the maximum value. We also consider the issue whether in the opposite direction with respect to the M31 galaxy the same effect induces a minimum in temperature in the *Planck's* maps of the sky. We find that the *Planck's* data at 100 GHz show an effect even larger than that expected.

Key words. Galaxies: general – Galaxies: individual (M31) – Galaxies: disks – Galaxies: halos

1. Introduction

Galactic disk rotation can be accurately investigated in the optical, infrared (IR) and radio bands and allows to infer important information, among others, about the dynamical mass content of galaxies (see e.g. Binney & Merrifield 1998). On the other hand, many ambiguities still exist about the main constituents of the galactic halos. The degree to which galactic halos rotate with respect to the disks is a particularly difficult task to be investigated, even for the closest galaxy to the Milky Way: M31 (Courteau et al. 2011). A novel approach in the study of the rotation of either the disk and halo of nearby galaxies (particularly the M31 galaxy) has been discussed in De Paolis et al. (2011). By using the 7-year WMAP data, a possible temperature asymmetry was found both in the M31 disk and halo along the direction of the M31 rotation, therefore reminiscent of a Doppler-induced effect. By adopting the geometry described in Fig. 1 in De Paolis et al. (2011), and extending the analysis up to about 20^0 ($\simeq 260$ kpc) around the M31 center, we found in the two opposite regions of the M31 disk a temperature difference of about $130 \mu\text{K}$, more or less the same in the W, V and Q WMAP bands. A similar effect was visible also towards the M31 halo up to about 120 kpc from the M31 center with a peak value of about $40 \mu\text{K}$. The robustness of that result was tested by considering 500 randomly distributed control fields and also simulating 500 sky map realizations from the best-fit power spectrum

constrained with BAO and H_0 (see De Paolis et al. 2011 for details). It turned out that the probability that the detected temperature asymmetry towards the M31 disk is due to a random fluctuation of the CMB signal is below about 2% while in the case of the M31 halo it is less than about 30%. Although the confidence level of the signal was not high with WMAP data, if estimated purely statistics, nevertheless we believed that the geometrical structure of the temperature asymmetry pointed towards a definite effect modulated by the rotation of the M31 disk and halo and suggested that with the *Planck* data it could be possible to definitely prove or disprove our conclusions. Indeed, the *Planck* satellite is about ten times more sensitive than the WMAP satellite and has an angular resolution about three times better: the *Planck* full width half maximum (FWHM) resolution ranges from $33.3'$ to $4.3'$ going from 30 GHz to 857 GHz, and its final sensitivity is in the range of $2 - 4.7 \mu\text{K/K}$ in terms of $\delta T/T$ for the Low Frequency Instrument (LFI), that is in the range $30 - 70$ GHz, and of $2 - 14 \mu\text{K/K}$ for the High Frequency Instrument (HFI) below 353 GHz (see e.g. Burigana et al. 2013 for a recent review on *Planck's* results). The aim of the present paper is therefore to analyze in detail the *Planck* data acquired during the first 15.4 months of observations towards both the disk and halo of the M31 galaxy. In addition, we also take the opportunity of investigating in some detail the recent issue raised in Rubin & Loeb (2014) about the kinetic Sunyaev-Zel'dovich effect from the diffuse hot gas in the Local Group, which happens to show up as a hot spot of a few degrees in size in the direction of M31 (where the free

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electron optical depth gets the maximum value). We also investigate whether in the opposite direction with respect to the M31 galaxy, the same effect induces a minimum in temperature in the *Planck*'s maps of the sky.

2. Planck analysis

Two instruments are present onboard the *Planck* satellite: the LFI (Bersanelli et al. 2010) covers the 30, 44, and 70 GHz bands by using amplifiers cooled to 20 K. The HFI (Lamarre et al. 2010) covers the 100, 143, 217, 353, 545, and 857 GHz bands with bolometers cooled down to 0.1 K. *Planck*'s sensitivity, angular resolution (from 3' to 5') and frequency coverage make it a powerful instrument for cosmology and both galactic and extragalactic astrophysics (Aghanim et al. 2012). In order to reveal the different contribution by the M31 disk and halo, the region of the sky around the M31 galaxy has been divided into several concentric circular areas as shown in Fig. 1 in De Paolis et al. (2011), to which we refer for further details. Here we only mention that the M1 region is the M31 south-east half-disk while the M2 region corresponds to the north-west half-disk. Since the M31 disk is rotating in the clock-wise direction we expect that the M1 region would be hotter than the M2 one. The mean temperature excess T_m in μK in each region has been obtained in each *Planck*'s band and is shown in Table 1 with the corresponding standard error (SE)¹, along with the number of pixels in each area.

2.1. Results for the M31 disk

As far as the M31 disk is concerned and as can be seen from the first four lines of Table 1 and Fig. 1, the M1 region turns out to be always hotter than the M2 region. For example, at 1.5° the M1 region is $67 \mu\text{K}$ hotter than the M2 region and even at 4° the M1 region is $38 \mu\text{K}$ hotter than M2. The upper panel of Fig. 1 clearly shows the temperature asymmetry profile in the two regions of the M31 disk. These profiles are in agreement with the results obtained previously by using the WMAP data (De Paolis et al. 2011).² Moreover, the shape of the two profiles is clearly mirror-like, as expected if the effect is due to a Doppler modulation induced by the M31 disk rotation. Indeed, the hotter (M1) region corresponds to the side of the M31 disk that rotates towards us. This mirror-like shape of the two regions of the M31 disk is also visible in the M31 thick HI disk obtained at 21 cm (Chemin et al. 2009; Corbelli et al. 2010). In order to test whether the temperature asymme-

try we see towards the M31 disk is real or can be explained as a random fluctuation of the CMB signal (that indeed is rather patchy) we adopt a different strategy with respect to that in De Paolis et al. (2011). We consider 360 control field regions with the same shape as the M1 and M2 regions and at the same latitude as M31 but at 1° longitude from each other. For each region we determine the excess temperature profile and calculate the average profile and the corresponding standard deviation. As can be easily observed by looking at Table 1 and at the bottom panel of Fig. 1, the M1 region for the 360 control fields is always cooler than the M2 region, exactly the opposite of what happens towards the M31 disk. Moreover, the M1 temperature towards M31 is always significantly larger with respect to the corresponding temperature of the control fields and even at 4° the effect is at $\simeq 4\sigma$. The same also holds for the M2 region: the 360 control fields have a temperature excess of $66 \pm 10 \mu\text{K}$ at 4° while the temperature of the M31 M2 region is always cooler, being $\simeq 32 \mu\text{K}$. The effect is therefore at $\simeq 3\sigma$ at 4° . We remark that we have conducted the same study in all the *Planck* bands and find that the results, presented for convenience only for the 100 GHz band in this paper, are comparable in each band.

2.2. Results for the M31 halo

Adopting the same geometry as in De Paolis et al. (2011), we have estimated the temperature excess in the *Planck*'s sky maps in the N1+S1 region (the region in the south-east of the M31 halo that is expected to be rotating moving towards the Milky way, if the M31 halo is rotating along the same axis of the disk) and in the N2+S2 region (the opposite region with respect to the rotation axis). As can be seen from Table 1 and Fig. 2 (upper panel), the N1+S1 region turns out to be hotter than the N2+S2 one at any galactocentric distance. The temperature difference peaks at about 4° (with a value about $38 \mu\text{K}$), but continues up to 10° (where it is still at $\simeq 13 \mu\text{K}$). Beyond about 12° the temperature asymmetry gets inverted and the N2+S2 region becomes hotter than the N1+S1 one, as a result of the intersection with the Milky Way disk that clearly shows up in the CMB maps. As for the M31 disk, also for the halo one can observe a kind of mirror symmetry between the N1+S1 and N2+S2 regions, although less pronounced with respect to the case of the M31 disk. We have also tested whether the measured temperature asymmetry is due to a random fluctuation of the CMB signal by considering 360 control fields with the same shape of the N1+S1 and N2+S2 regions at the same latitude of M31 but at different longitude values (the control fields are equally spaced at one degree distance each other in longitude). As one can see from the bottom panel of Fig. 2, also for the halo regions (as for the M31 disk) the temperature asymmetry in the 360 control fields clearly has an opposite behavior with respect to the profiles towards M31 and the N1+S1 regions are always cooler than the N2+S2 ones (bottom panel in Fig. 2). This effect is clearly due to the presence of the Milky Way disk in the CMB sky maps which makes the N2+S2 regions generally hotter than the N1+S1 regions. It can be easily observed by comparing the temperature asymmetry profile of the M31 halo with that of the control fields that the N1+S1 region of M31 is always hotter than the control fields profile (with a confidence level of about 4σ at 4° and 2.7σ at 10°) while the N1+S1 region is cooler than the control fields profile

¹ The standard error given in the fourth column is calculated as the standard deviation of the excess temperature distribution divided by the square root of the pixel number in each region. To possibly enable the comparison with the previous WMAP data analysis De Paolis et al. (2011), here we use *Planck*'s 100 GHz data and have verified that, within the errors, the sigma values calculated in that way are consistent with those evaluated by using the covariance matrix obtained by a best fitting procedure with a Gaussian to the same distribution. In the last column we give the average excess temperature for 360 control fields with the usual standard deviation.

² The absence of foreground reduced *Planck*'s maps, that where instead available for WMAP maps, makes ambiguous the comparison between real and simulated data. The strategy adopted here of using 360 control fields in the *Planck*'s maps gives more reliable results.

R, deg, kpc	Region	N, pix	$T_m \pm SE$	$T_m \pm \sigma$ for 360 control fields
1.5, 19.5	M1	4213	115.4 ± 2.0	44.0 ± 5.0
1.5, 19.5	M2	4182	48.2 ± 2.0	50.0 ± 6.0
4.0, 51.9	M1	29076	70.1 ± 0.9	41.0 ± 7.0
4.0, 51.9	M2	28983	32.0 ± 0.9	66.0 ± 10.0
4.0, 51.9	N1+S1	27957	70.0 ± 1.0	41.0 ± 7.0
4.0, 51.9	N2+S2	27874	32.2 ± 1.0	66.0 ± 9.0
10.0, 131.2	N1+S1	158752	65.0 ± 0.2	43.0 ± 8.0
10.0, 131.2	N2+S2	158720	52.2 ± 0.2	73.0 ± 10.0
4.0, 51.9	M31	61306	50.1 ± 0.3	44.6 ± 1.6
4.0, 51.9	anti M31	61306	-6.2 ± 0.3	24.8 ± 1.0

Table 1. Temperature excess in the M31 regions. The radius of the considered annulus is given in degrees and in kpc in the first column; the value of 744 kpc (Vilardell et al. 2010) is adopted for the distance to M31. The second column indicates the considered region. In the third column the numbers of pixels in each region are given. The fourth column show the CMB mean temperature of each region (in μK) in the 100 GHz *Planck* map with the corresponding standard error (SE) while the last column gives the average temperature excess in the 360 control fields with the standard deviation (see text for details).

(with a confidence level of about 3.7σ at 4° and 2.1σ at 10°). We can therefore conclude that the probability that the asymmetry effect towards the M31 halo at 100 GHz is due to a random fluctuation of the CMB signal is well below 1%. We also point out that we have verified that the temperature asymmetry towards the M31 halo vanishes if the adopted geometry is rigidly rotated by an angle larger than about 10° with respect to the assumed M31 rotation axis, thus giving a further indication that the asymmetric halo temperature is a genuine effect due to the halo rotation and not simply a random fluctuation of the CMB signal.

2.3. The Local Group hot gas effect

Recently Rubin & Loeb (2014) raised an interesting issue related to the kinetic Sunyaev-Zel'dovich effect from the diffuse hot gas in the Local Group. In fact, since the Local Group moves with respect to the CMB (Rauzy & Gurzadyan 1998), its hot gas halo component should imprint a non-primordial temperature shift in the CMB signal. The expected effect should show up as a hot spot of a few degree in size in the direction of the M31 galaxy, which happen to be opposite with respect to the center of the Local Group. In fact, due to geometrical consideration, the free electron optical depth gets the maximum value just towards the M31 galaxy. On the other hand, in the opposite direction with respect to the M31 galaxy, the same effect should induce a minimum in temperature in the *Planck*'s maps of the sky. We have investigated this issue by looking at the *Planck*'s sky map at 100 GHz and find (see the last two lines at the bottom of Table 1) that the mean temperature excess in a 4° circle towards the M31 galaxy is $\simeq 50.1 \pm 0.3 \mu\text{K}$, therefore consistently hotter with respect to the average temperature in the southern hemisphere of the sky ($\simeq 40.5 \mu\text{K}$)³. Towards the anti M31 direction ($l = 301.17^\circ$, $b = 21.57^\circ$) we find, in a 4° circle, a temperature excess of $-6.2 \pm 0.3 \mu\text{K}$, to be compared to the average temperature in the northern hemisphere of the sky of about $27.5 \mu\text{K}$. As for the M31 disk and halo, we also consider 360 control fields – this time randomly extracted around (within 20°) either the M31 and

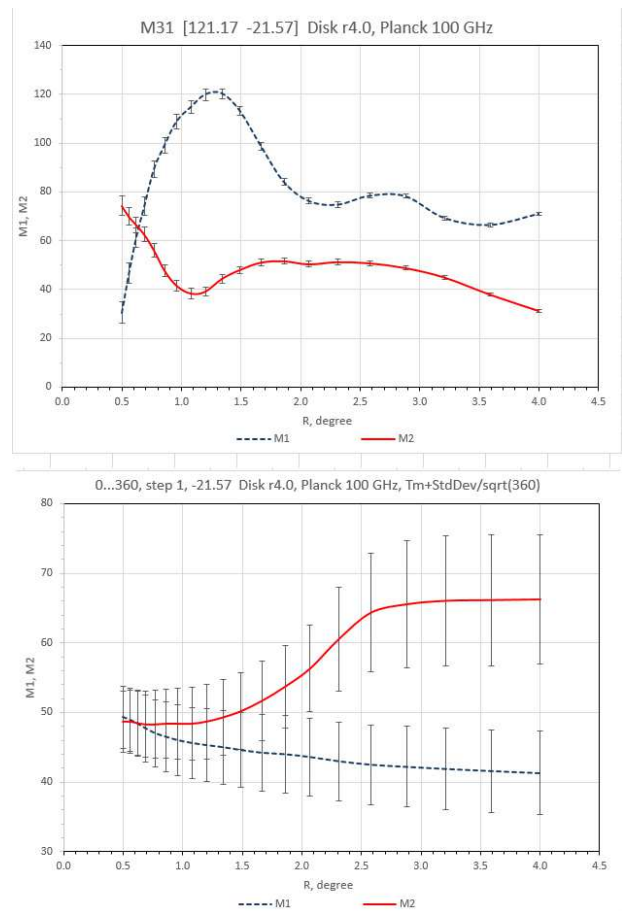


Fig. 1. Upper panel: the excess temperature profiles (in μK) for the M1 and M2 regions of the M31 disk. Bottom panel: temperature profiles (in μK) for 360 regions equally spaced at one degree distance each other in longitude and at the same latitude as M31.

anti M31 directions. The obtained results are shown in the two lines at the bottom of Table 1 and, as one can see, the 4° circle towards M31 is hotter than the control fields at about 3σ . Towards the anti M31 direction, the 4° circle is clearly much cooler with respect to the control fields (at about 29σ). The effect predicted towards the M31 galaxy by Rubin & Loeb (2014) was of a few μK and no mention-

³ The mean temperature in both the southern and northern hemispheres of the sky has been evaluated after the equatorial region, affected by the Milky Way emission, has been subtracted

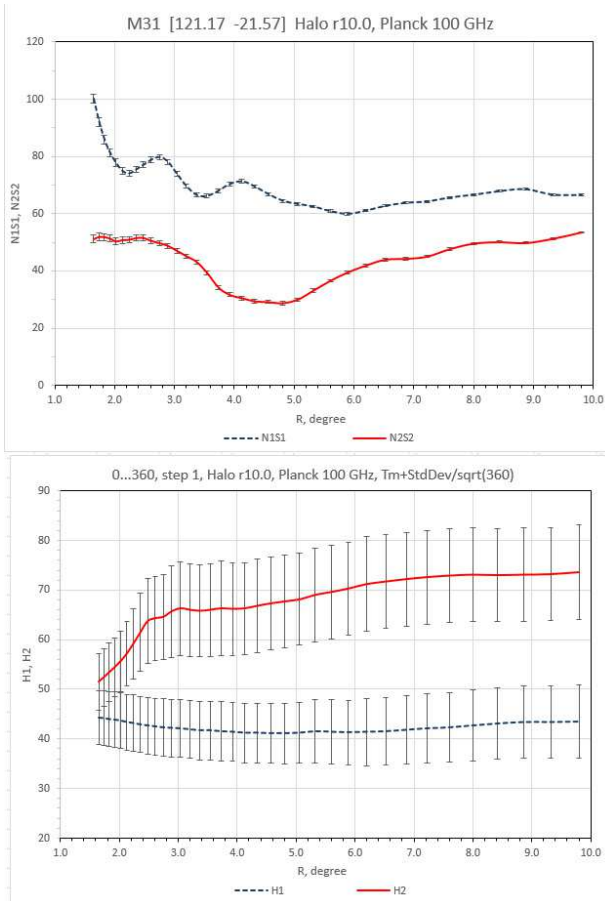


Fig. 2. The same as for Fig. 1 but for the M31 halo (upper panel) and for 360 regions at one degree longitudinal distance each other ($H=N+S$, bottom panel).

ing about the possible existence of a cold spot in the anti M31 direction was present there. From the discussion above it is clear that the observed temperature difference (about $56 \mu\text{K}$) between the M31 and anti M31 directions cannot be explained as a random fluctuation of the CMB signal and should therefore arise from two main contributions: the kinetic Sunyaev-Zel'dovich effect and the presence of hot gas in the M31 halo, with a density larger than the average hot gas density in the Local Group. The presence of this hot gas halo component in the M31 halo, as predicted in De Paolis et al. (1995), might be able to explain both the CMB temperature increase towards the M31 galaxy and, if it rotates around the same rotation axis as the M31 disk, the temperature shifts between the two sides of the M31 halo as discussed above.

3. Conclusions

Galactic halos are relatively less studied than galactic disks and there are still many ambiguities not only in the main halo constituents, but also with respect to the degree to which galactic halos rotates with respect to the disks (Courteau et al. 2011; Deason, Belokurov & Evans 2011). Actually, the rotation of the galactic halos is clearly related to the formation scenario of galaxies. In the standard collapse model (see e.g. Eggen, Lynden-Bell & Sandage 1962) both the halo and disk derive from the same population and the rotation of the outer halo should be, in this case, aligned

with the disk angular momentum. On the contrary, in a hierarchical formation scenario, structures reaching later the outer halo should be less connected to the disk. Therefore, it is evident that information on the galactic halo rotation provides key insights about the formation history of galaxies. It is also well known that the M31 disk rotates with a speed of about 250 km s^{-1} and this has been clearly shown also by the velocity maps obtained from radio measurements (Chemin et al. 2009; Corbelli et al. 2010). These maps look very similar to what we find in the *Planck* data towards the M31 disk. In the previous Section we have also shown that *Planck*'s data show the existence of a temperature asymmetry with respect to the disk-halo rotation axis, up to a galactocentric distance of about 130 kpc and with a peak temperature contrast of about $40 \mu\text{K}$. We remark that, until now, the only evidence of the M31 halo rotation was put forward by the analysis of the dwarf galaxies orbiting M31 (Ibata et al. 2013).

In all generality, five possibilities may be considered in order to explain the effects discussed in Sections 2.1-2.3: (i) free-free emission; (ii) synchrotron emission; (iii) anomalous microwave emission (AME) from dust grains; (iv) kinetic Sunyaev-Zel'dovich (SZ) effect; (v) cold gas clouds populating the M31 halo. A detailed study of their contribution using all the *Planck*'s bands to constrain the model parameters and the relative weight of these five models will be published elsewhere. Here, we only note that effects (i) – (iii) give a signal with a rather strong dependence on the wavelength, while (iv) and (v) are almost independent of the observation band in the microwave regime and to first approximation could provide the main contribution to the observed effect. Thus, our investigation shows the power of CMB to trace, along with the clusters of galaxies via Sunyaev-Zeldovich effect and the large scale voids (e.g. Gurzadyan & Kocharyan 2009), also the individual galactic halos.

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References

- Aghanim, N., Arnaud, M., Ashdown, M. et al. 2012, *A&A*, 543, id. A102
- Bersanelli, M., Mandolesi, N., Butler, R. C., et al. 2010, *A&A*, 520, A4
- Binney, J. & Merrifield, M. *Galactic Astronomy*, Princeton Series in Astrophysics (1998)
- Burigana, C., Davies, R.D., De Bernardis, P. et al. 2013, *IJMPD*, 22, id. 1330011
- Chemin, L., Carignan, C. & Foster, T. 2009, *ApJ*, 705, 1395
- Górski, K. M., Hivon, E., Banday, A. J. et al. 2005, *ApJ*, 622, 759
- Gurzadyan, V.G. & Kocharyan, A.A., 2009, *A&A*, 493, L61
- Ibata, R. A. et al. 2013, *Nature*, 493, 62
- Lamarre, J.M., Puget, J.L., Ade, P. A. R. et al. 2010, *A&A*, 520, id. A9
- Rauzy, S. & Gurzadyan, V.G. 1998, *MNRAS*, 298, 114
- Rubin, D. and Loeb, A. 2014, *arXiv:1311.5255v2*
- Vilardell, F. et al. 2010, *A&A*, 509, 70